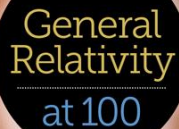


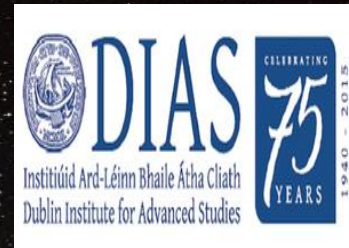
A new perspective on Einstein's cosmology

Cormac O'RaiFeartaigh

Waterford Institute of Technology
Dublin Institute for Advanced Studies



General
Relativity
at 100



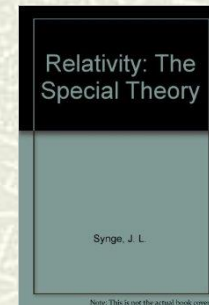
Brendan McCann, Michael O'Keeffe (Waterford IT)
Simon Mitton (Cambridge Univ.), Werner Nahm (Dublin IAS)

Dublin Institute for Advanced Studies

75 years of DIAS

Founded in 1940 (de Valera)

Modelled on Princeton Research Institute (IAS)

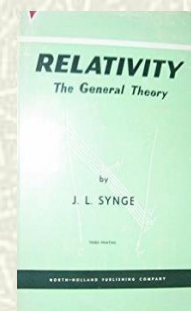
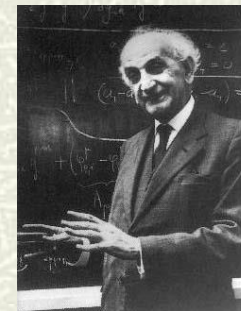


Two schools

School of Theoretical Physics

School of Celtic Languages

Only pen and paper required

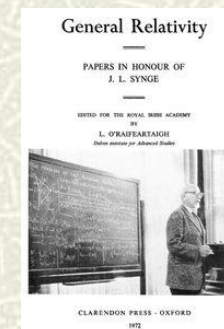


Major centre for relativity

Erwin Schrödinger, J.L. Synge,

K. Lanczos, L. O'Raifeartaigh

Pirani, Bertotti,...



Overview

Introduction

Einstein's cosmological considerations 1915-1930

1. Einstein's cosmic model of 1931

Some new observations

2. Einstein's steady-state model

Unfinished symphony

3. Einstein's cosmology review of 1933

Further development of the Einstein-de Sitter model

Conclusions

$$D = \frac{1}{c} \frac{1}{L} \frac{dL}{dt} = \frac{1}{c} \frac{1}{P} \frac{dP}{dt}$$
$$D^2 = \frac{1}{P^2} \frac{P_0 - P}{P} \sim \frac{1}{P^2} \quad (1a)$$
$$D^2 = \frac{K \rho}{3} \frac{P_0 - P}{P} \sim \frac{1}{3} K \rho \quad (2a)$$
$$D^2 \sim 10^{-53}$$
$$\rho \sim 10^{-26}$$
$$P \sim 10^8 \text{ (g. J.)}$$
$$t \sim 10^{10} \text{ (10}^{11} \text{)} \text{ } \}$$

The Oxford blackboard (1931)



Einstein's cosmology (1915-1930)

The general theory of relativity (1915, 1916)

A relation between spacetime and matter

Gravity = curvature of spacetime

$$ds^2 = -c^2 dt^2 + dx^2 + dy^2 + dz^2$$

$$ds^2 = g_{\mu\nu} dx^\mu dx^\nu$$



Principles

Principle of equivalence

Mach's principle: no inertia relative to space

Matter has a primary role, space has a derived one

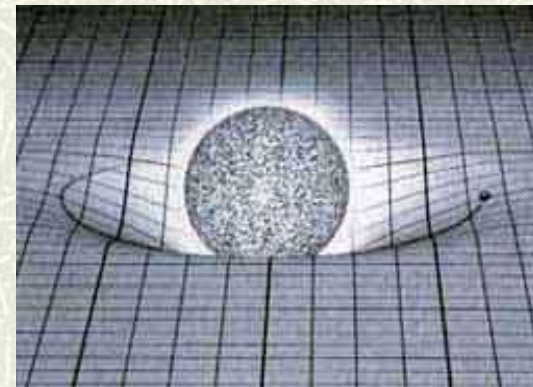
$$R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R = -\kappa T_{\mu\nu}$$

A burning question: a consistent model of the Cosmos?

"Can relativity be followed through to the finish?" (1918)

Cosmological considerations (1917)

Result of long deliberations: "A rough and winding road"



Einstein's model of the Static Universe

Apply general relativity to the Universe (1917)

Ultimate test for new theory of gravitation

Assumptions

Static universe (small velocities of the stars)

Mach's principle (metric tensor to vanish at infinity)

Isotropy and homogeneity (simplicity)

Boundary problem

A cosmos of closed curvature

No consistent solution

Field equations modified!

Additional term in GFE (1916)

Radius and density defined by λ

$$R_{\mu\nu} - \frac{1}{2}g_{\mu\nu}R = -\kappa T_{\mu\nu}$$

$$R_{\mu\nu} - \frac{1}{2}g_{\mu\nu}R - \lambda g_{\mu\nu} = -\kappa T_{\mu\nu}$$

Doc. 43

Cosmological Considerations in the General Theory of Relativity

This translation by W. Perrett and G. B. Jeffery is reprinted from H. A. Lorentz et al., *The Principle of Relativity* (Dover, 1952), pp. 175–188.

It is well known that Poisson's equation $\nabla^2\phi = 4\pi K\rho$ (1) in combination with the equations of motion of a material point is not as yet a perfect substitute for Newton's theory of action at a distance. There is still to be taken into account the condition that at spatial infinity the potential ϕ tends



$$\lambda = \frac{\kappa\rho}{2} = \frac{1}{R^2}$$

Some key quotes (Einstein 1917)

“In a consistent theory of relativity, there can be no inertia relative to “space”, but only an inertia of masses relative to one another”

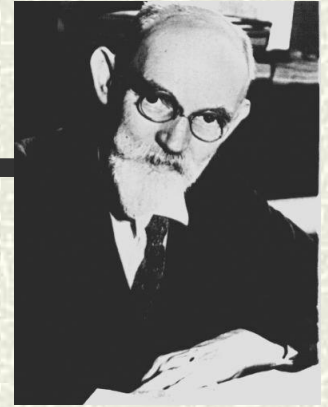
“I have not succeeding in formulating boundary conditions for spatial infinity. Nevertheless, there is still a way out...for if it were possible to regard the universe as a continuum which is finite (closed) with respect to its spatial dimensions, we should have no need at all of any such boundary conditions”

“The most important fact that we draw from experience as to the distribution of matter is that the relative velocities of the stars are very small compared with the velocity of light..... There is a system of reference relative to which matter may be looked upon as being permanently at rest ”

“However, the system of equations ..allows a readily suggested extension which is compatible with the relativity postulate... For on the left hand side of the field equation...we may add the fundamental tensor $g_{\mu\nu}$, multiplied by a universal constant , $-\lambda$, at present unknown, without destroying the general covariance ”

Schroedinger’s comment (1918): Einstein’s response (1918)

Einstein vs de Sitter



Alternative solution of the GFE

A universe empty of matter (1917)

Solution B

Cosmic constant proportional to curvature of space

$$\lambda = 3/R$$

Disliked by Einstein

Problems with singularities? (1918)

Conflict with Mach's principle

The de Sitter confusion

Static or non-static - a matter of co-ordinates?

Weyl, Lanczos, Lemaître

[p. 270] 5. "Critical Comment on a Solution of the Gravitational Field Equations Given by Mr. De Sitter"

[Einstein 1918c]

SUBMITTED 7 March 1918

PUBLISHED 21 March 1918

IN: *Königlich Preußische Akademie der Wissenschaften (Berlin). Sitzungsberichte* (1918): 270-272.

[1] Herr De Sitter, to whom we owe deeply probing investigations into the field of the general theory of relativity, has recently given a solution for the equations of gravitation which, in his opinion, could possibly represent the metric structure of the universe. However, it appears to me that one can raise a grave argument against the admissibility of this solution, which shall be presented in the following.

The De Sitter solution of the field equations

$$G_{\mu\nu} - \lambda g_{\mu\nu} = -\kappa T_{\mu\nu} + \frac{1}{2}g_{\mu\nu}\kappa T \quad (1)$$

is

Prediction of redshifts – Slipher effect?

Einstein vs Friedman



Alexander Friedman
(1888 -1925)

Allow time-varying solutions (1922)

Assume homogeneity, isotropy, positive curvature

Two independent differential equations from GFE

$$\frac{3R'^2}{R^2} + \frac{3c^2}{R^2} - \lambda = \kappa c^2 \rho,$$

$$\frac{R''}{R^2} + \frac{2RR''}{R^2} + \frac{c^2}{R^2} - \lambda = 0.$$

Evolving universes

Density of matter varies over time

$$\frac{1}{c^2} \left(\frac{dR}{dt} \right)^2 = \frac{A - R + \frac{\lambda}{3c^2} R^3}{R}$$

Overlooked by community

Considered 'suspicious' by Einstein

Mathematical correction, later retracted

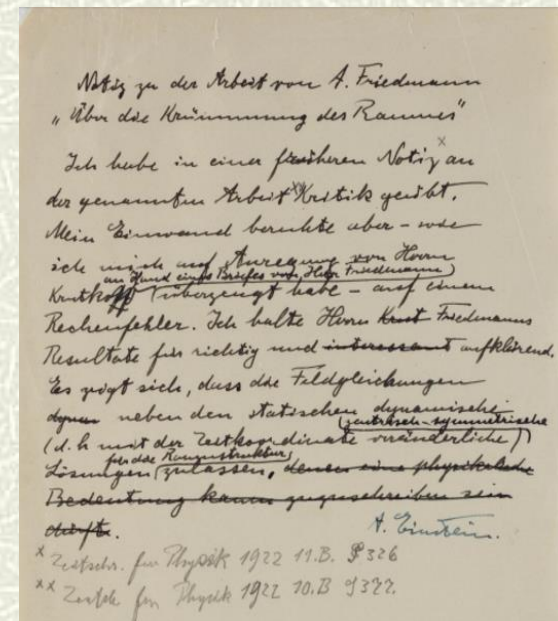
~~"To this a physical reality can hardly be ascribed"~~

$$t = \frac{1}{c} \int_a^R \sqrt{\frac{x}{A - x + \frac{\lambda}{3c^2} x}} dx + B$$

Negative spatial curvature (1924)

Cosmic evolution, geometry depends on matter

Overlooked by community



Einstein vs Lemaître (1927)



Fr Georges Lemaître

Expanding model of the Universe from GR

Similar but not identical to Friedman 1922

Starts from static Einstein Universe

$$3\frac{R'^2}{R^2} + \frac{3}{R^2} = \lambda + \kappa\rho$$

$$2\frac{R''}{R} + \frac{R'^2}{R^2} + \frac{1}{R^2} = \lambda - \kappa\rho$$

Redshifts of galaxies = expansion of metric?

Redshifts from Slipher, distances from Hubble

$$H = 585 \text{ kms}^{-1}\text{Mpc}^{-1}$$

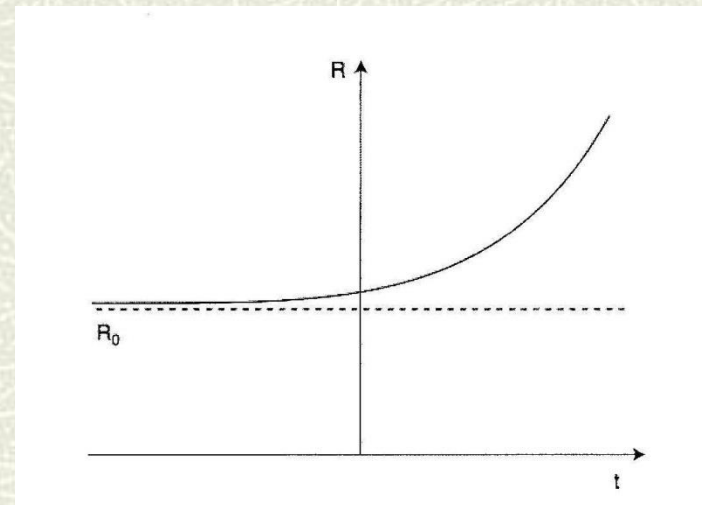
Ignored by community

Belgian journal (in French)

Rejected by Einstein: "Votre physique est abominable"

Lemaître's recollection (1958)

"Einstein not up-to-date with astronomy"



Hubble's law (1929)



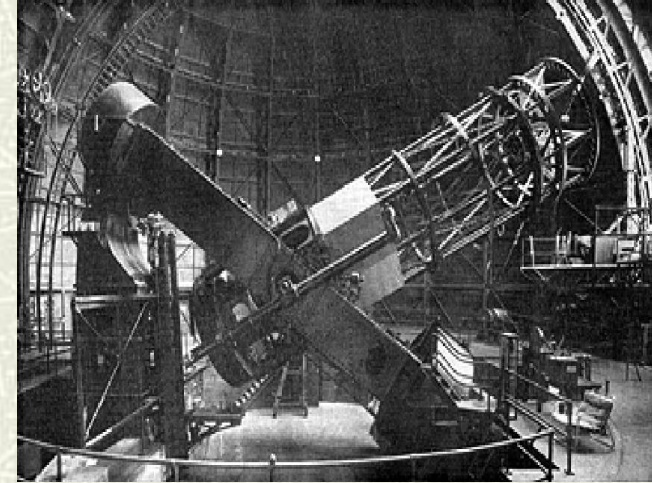
■ A redshift/distance relation for the nebulae?

Motivation: establishing distances of all nebulae

■ Linear relation (Hubble, 1929)

20 redshifts from Slipher: not acknowledged

Most important data point not shown (8Mpc, 4000 km/s)

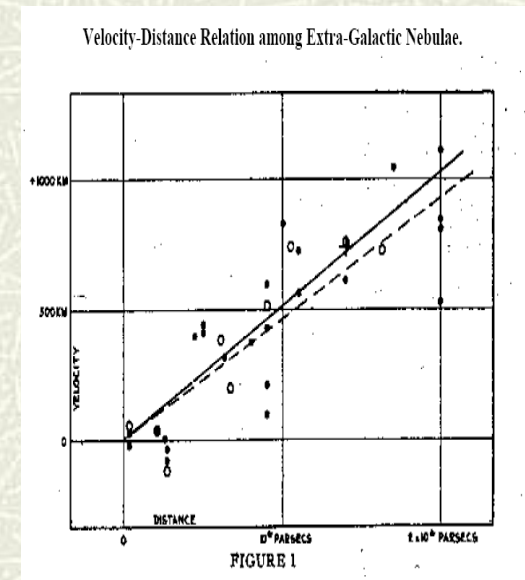


■ Landmark result in astronomy

$H = 500 \text{ kms}^{-1} \text{ Mpc}^{-1}$

■ Not the expanding universe!

Astronomy, not cosmology



The watershed

- **RAS meeting (1930)**

*If redshifts are velocities, and if effect is non-local
Hubble's law = expansion of space? (Edd., de Sitter)*

- **Dynamic model required**

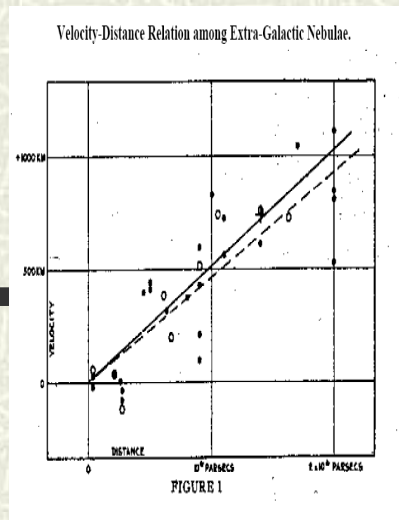
"Motion in Einstein's model or matter in de Sitter's?"

- **Lemaître's intervention**

*1927 expanding model republished in English (1931)
Observational section omitted (rightly)*

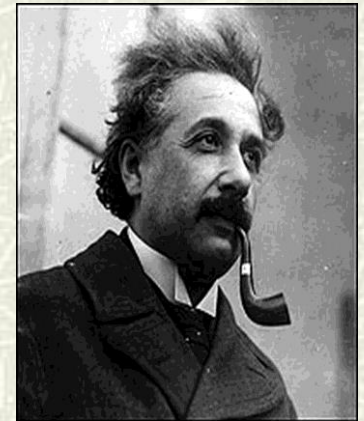
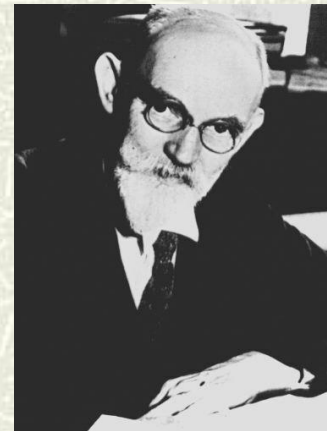
- **Lemaître model circulated**

*Time-varying radius, density of matter
Friedman 1922 models become known
Positive curvature*



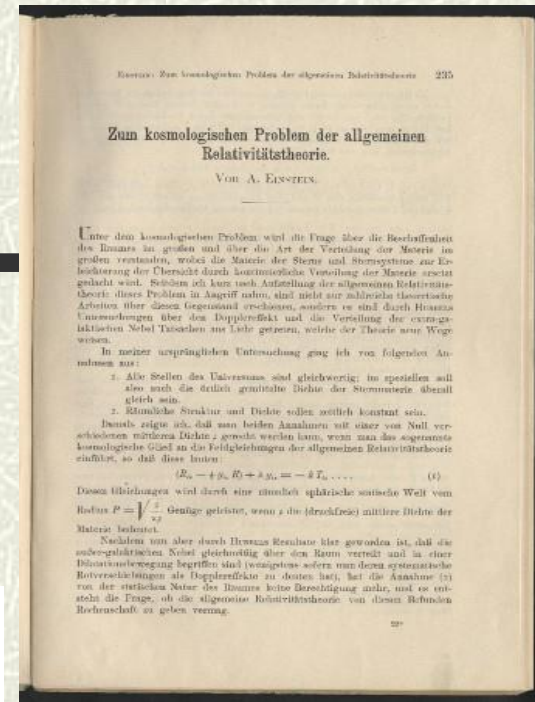
Models of the evolving universe (1930 -)

- **Tolman (1930, 31)**
Expansion caused by annihilation of matter ?
- **Eddington (1930, 31)**
On the instability of the Einstein universe
Expansion caused by condensation?
- **de Sitter (1930, 31)**
Variety of expanding models
- **Heckmann (1931,32)**
Spatial curvature (not translated)
- **Einstein (1931, 32) JMK**
Friedman-Einstein model $\lambda = 0, k = 1$
Einstein-de Sitter model $\lambda = 0, k = 0$



If redshifts represent expansion...
If effect is global...

Einstein's 1931 model (F-E)



■ Einstein's 'first' model of the Expanding Universe

Occasionally cited, rarely read (not translated)

$$\frac{3P'^2}{P^2} + \frac{3c^2}{P^2} - \lambda = \kappa c^2 \rho.$$

■ Adopts Friedman 1922 model

Time-varying, closed universe: $k = 1$

$$\frac{P'^2}{P^2} + \frac{2P''}{P} + \frac{c^2}{P^2} - \lambda = 0$$

■ Set cosmic constant to zero

Instability of static solution

Hubble's observations

$$D = \frac{1}{P} \frac{dP}{dt} \cdot \frac{1}{c}$$

$$D^2 = \frac{1}{P^2} \frac{P_0 - P}{P}$$

$$P \sim \frac{1}{D}$$

$$\left(\frac{dP}{dt}\right)^2 = c^2 \frac{P_0 - P}{P}$$

■ Extraction of cosmic parameters!

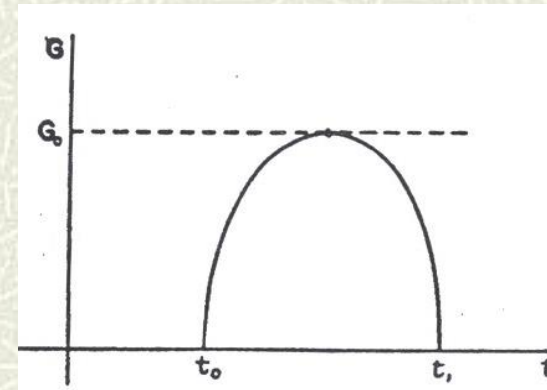
$$P \sim 10^8 \text{ yr} : \rho \sim 10^{-26} \text{ g/cm}^3$$

$t \sim 10^{10}$ yr : conflict with astrophysics

Attributed to simplifying assumptions (homogeneity)

$$D^2 = \frac{1}{3} \kappa \rho \frac{P_0 - P}{P}$$

$$D^2 \sim \kappa \rho$$



Einstein's 1931 model revisited

First translation into English

O'Raifeartaigh and McCann 2014

$$D = \frac{1}{P} \frac{dP}{dt} \cdot \frac{1}{c}$$

$$D^2 = \frac{1}{P^2} \frac{P_0 - P}{P}$$

$$P \sim \frac{1}{D}$$

Anomalies in calculations of radius and density

Einstein: $P \sim 10^8$ yr, $\rho \sim 10^{-26}$ g/cm³, $t \sim 10^{10}$ yr

We get: $P \sim 10^9$ yr, $\rho \sim 10^{-28}$ g/cm³, $t \sim 10^9$ yr

$$D^2 = \frac{1}{3} \kappa \rho \frac{P_0 - P}{P}$$

$$D^2 \sim \kappa \rho$$

Source of error?

Oxford blackboard: $D^2 \sim 10^{-53}$ cm⁻² should be 10^{-55} cm⁻²

Time miscalculation $t \sim 10^{10}$ yr (should be 10^9 yr)

Non-trivial error: misses conflict with radioactivity

*Oxford lecture
(May 1931)*

Not a cyclic model

"Model fails at $P = 0$ "

Contrary to what is usually stated

$$D = \frac{1}{c} \frac{1}{P} \frac{dP}{dt}$$

$$D^2 = \frac{1}{P^2} \frac{P_0 - P}{P} \sim \frac{1}{P^2} \quad (1a)$$

$$D^2 = \frac{1}{3} \kappa \rho \frac{P_0 - P}{P} \sim \kappa \rho \quad (2a)$$

$$D^2 \sim 10^{-53}$$

$$\rho \sim 10^{-26}$$

$$P \sim 10^8 \text{ yr}$$

$$t \sim 10^{10} (10^9) \text{ yr}$$

Some key quotes (Einstein 1931)

“The cosmological problem is understood to concern the question of the nature of space and the manner of the distribution of matter on a large scale, where the material of the stars and stellar systems is assumed for simplicity to be replaced by a continuous distribution of matter.”

“Now that it has become clear from Hubbel’s results that the extra-galactic nebulae are uniformly distributed throughout space and are in dilatory motion (at least if their systematic redshifts are to be interpreted as Doppler effects), assumption (2) concerning the static nature of space has no longer any justification....”

“Several investigators have attempted to account for the new facts by means of a spherical space whose radius P is variable over time. The first to try this approach, uninfluenced by observations, was A. Friedman,¹ on whose calculations I base the following remarks. ”

“However, the greatest difficulty with the whole approach... is that according to (2 a), the elapsed time since $P = 0$ comes out at only about 10^{10} years. One can seek to escape this difficulty by noting that the inhomogeneity of the distribution of stellar material makes our approximate treatment illusory.”

Albert Einstein



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Albert Einstein



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Archival Call Number: 2-112
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 Main Author: Einstein, Albert (Author)
 Language: German
 Archival Location: Albert Einstein Archives, The Hebrew University of Jerusalem, Israel
 Number of Pages: 4



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By: Einstein, Albert (Author)
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Zum kosmologischen Problem der Allgemeinen Relativitätstheorie.

By: Einstein, Albert (Author)
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Zum kosmologischen Problem der allgemeinen Relativitätstheorie.

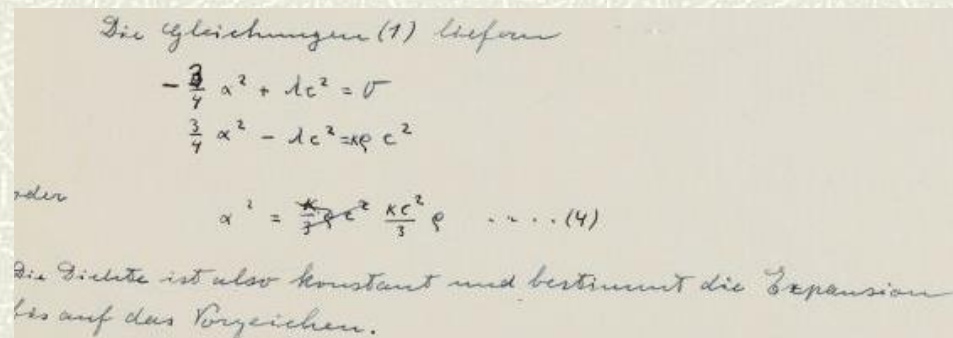
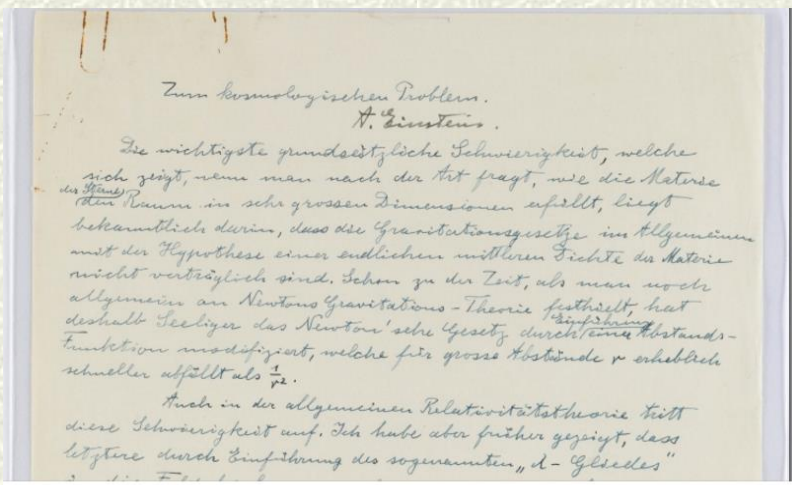
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Beiträge zum kosmologischen Problem.

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3. Zum kosmologischen Problem der Allgemeinen Relativitätstheorie. (4-53)
4. Kurze Zusammenfassung. (4-54)



“The density is thus constant and determines the expansion”

2. Einstein's steady-state model

Unpublished manuscript

Archived as draft of F-E model (1931)

Similar title, opening to F-E model

Something different

Cosmological constant

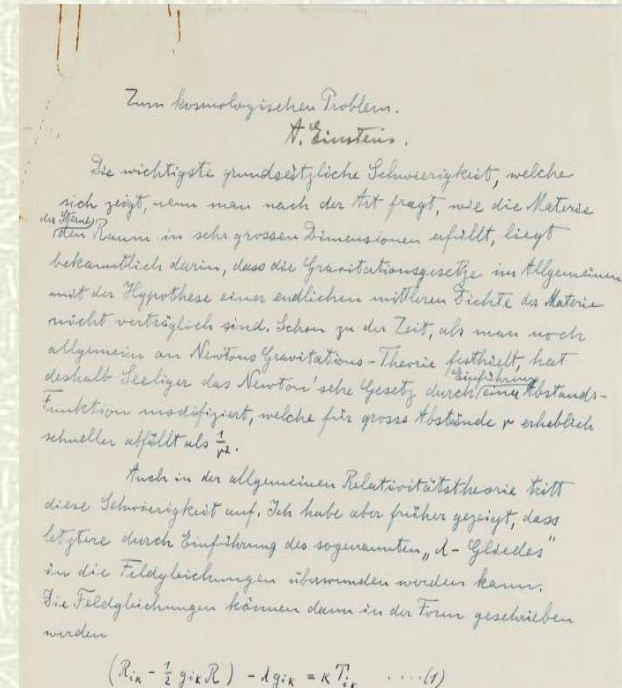
"The density is thus constant and determines the expansion"

Steady-state model of the Expanding Universe

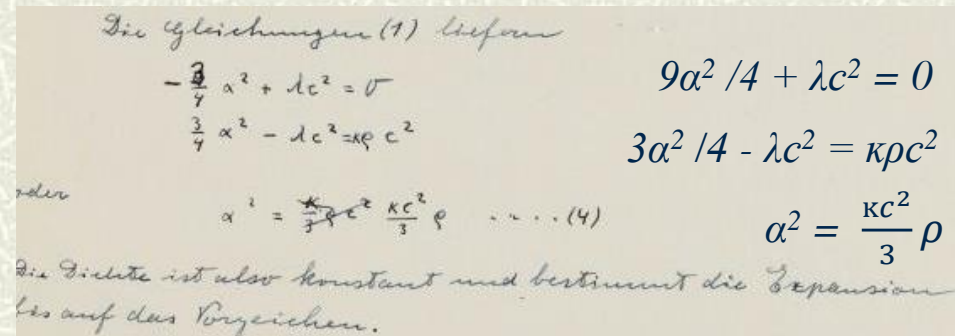
Anticipates Hoyle solution

Written in early 1931

Fatal flaw: abandoned



$$(R_{ik} - \frac{1}{2} g_{ik} R) - \lambda g_{ik} = \kappa T_{ik} \dots (1)$$



Einstein's steady-state model (Jan 31)

Problem with evolving models

“De Sitter and Tolman have already shown that there are solutions to equations (1) that can account for these [Hubbel's] observations. However the difficulty arose that the theory unvaryingly led to a beginning in time about 10^{10} – 10^{11} years ago, which for various reasons seemed unacceptable.”

New solution

“In what follows, I wish to draw attention to a solution to equation (1) that can account for Hubbel's facts, and in which the density is constant over time..

If one considers a physically bounded volume, particles of matter will be continually leaving it. For the density to remain constant, new particles of matter must be continually formed within that volume from space “

Mechanism

“The conservation law is preserved in that, by setting the λ -term, space itself is not empty of energy; its validity is well known to be guaranteed by equations (1).”

An abandoned model

✦ Correct geometry

de Sitter metric

✦ Simultaneous equations

Eliminate λ

Relation between α^2 and ρ

✦ Einstein's crossroads

Null solution on revision

Tolman? (Nussbaumer 2014)

Declined to amend GFE

✦ Evolving models

Less contrived: set $\lambda = 0$

Im Nachfolgenden will ich auf eine Lösung der Gleichung (1) aufmerktsamer machen, welche Hubble's Thatsachen gerecht wird, und in welcher die Dichte zeitlich konstant ist. Diese Lösung ist zwar in dem allgemeinen Schema Tolman's enthalten, scheint aber bisher nicht in Betracht gezogen worden zu sein.

1. Ich setze an

$$ds^2 = -e^{\alpha t} (dx_1^2 + dx_2^2 + dx_3^2) + c^2 dt^2 \dots (3)$$

Die Gleichungen (1) liefern

$$-\frac{3}{4} \alpha^2 + \lambda c^2 = 0 \qquad 9\alpha^2 / 4 + \lambda c^2 = 0$$




$$\frac{3}{4} \alpha^2 - \lambda c^2 = \kappa \rho c^2 \qquad 3\alpha^2 / 4 - \lambda c^2 = \kappa \rho c^2$$

oder

$$\alpha^2 = \frac{\kappa}{3} \rho c^2 \dots (4) \qquad \alpha^2 = \frac{\kappa c^2}{3} \rho$$

Die Dichte ist also konstant und bestimmt die Expansion bis auf das Vorzeichen.

Der Erhaltungssatz bleibt dadurch unvabert, dass bei Setzung des λ -Gledes der Raum selbst nicht energetisch leer ist; seine Ubertung wird bekanntlich durch die Gleichungen (1) gewahrleistet.




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Einstein's lost theory uncovered

Physicist explored the idea of a steady-state Universe in 1931.

Daive Castelvechi

24 February 2014

Physics » Nature   Email  Print


Einstein's Lost Theory Uncovered

The famous physicist explored the idea of a steady-state universe in 1931

nature

Feb 25, 2014 | By **Daive Castelvechi** and Nature magazine

A manuscript that lay unnoticed by scientists for decades has revealed that **Albert Einstein** once dabbled with an



New Discovery Reveals Einstein Tried To Devise A Steady State Model Of The Universe

www.irishtimes.com/news/science/wit-researchers-discover-lost-einstein-model-of-universe-1.1713487

THE IRISH TIMES **Science** Monday, March 10, 2014

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2 comments, 2 called-out [+ Comment Now](#) [+ Follow Comments](#)

Almost 20 years before the late Fred Hoyle and his colleagues devised the [Steady State Theory](#), Albert Einstein toyed with a similar idea: that the universe was eternal, expanding outward with a consistent input of spontaneously generating matter.

An Irish physicist came across the paper last year and could hardly believe. According to this week's article in [Nature](#),

model of the universe very different to today's [Big Bang](#) Theory.

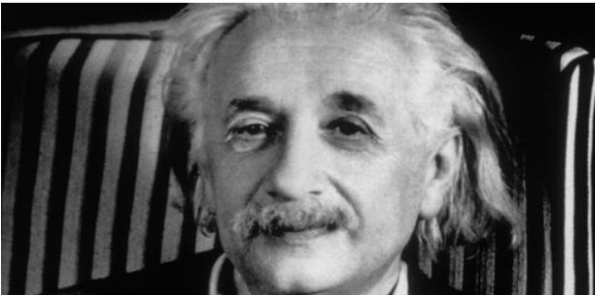
The manuscript, which hadn't been referred to by scientists for decades,



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WIT researchers discover 'lost' Einstein model of universe

Scientists uncovered misfiled papers while searching Jerusalem university's online archive




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- 08:25 Flannery faces call from all parties to attend PAC

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The way back isn't so simple



A useful find

New perspective on steady-state theory (1950s)

Logical idea: not a crank theory

Tolman, Schroedinger, Mimura : considered steady-state universe

Insight into scientific progress

Unsuccessful theories important in the development of science

Links with modern cosmology

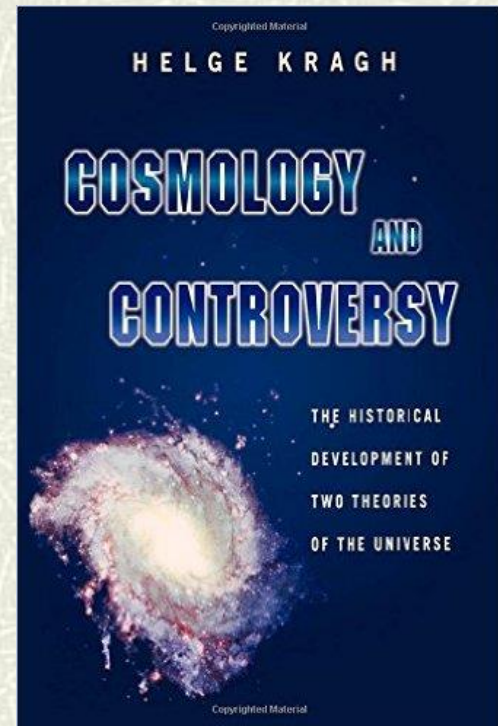
Creation energy and λ : dark energy

de Sitter metric: cosmic inflation

Insight into Einstein's cosmology

Turns to evolving models rather than introduce new term to GFE

Pragmatic approach: F-E model



3. Einstein-de Sitter model (1932)

Curvature not a given in dynamic models (Heckmann)

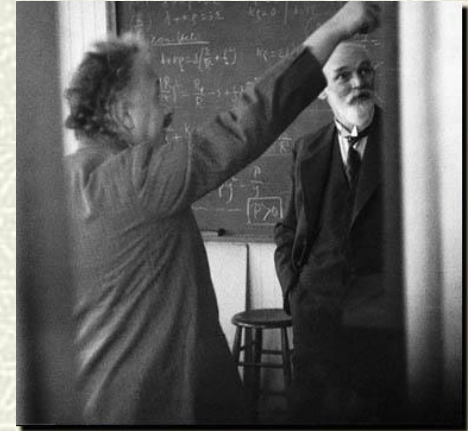
Not observed empirically

Remove spatial curvature (Occam's razor)

$$ds^2 = -R^2(dx^2 + dy^2 + dz^2) + c^2dt^2$$

$$\frac{1}{R^2} \left(\frac{dR}{cdt} \right)^2 = \frac{1}{3} \kappa \rho.$$

$$h^2 = \frac{1}{3} \kappa \rho$$



Simplest Friedman model

Time-varying universe with $\lambda = 0, k = 0, p = 0$

Estimate of density : $\rho = 10^{-28} \text{ g/cm}^3$

Important hypothetical cosmos: critical case

Becomes standard model

Despite high density of matter, age problem

Time evolution not considered

Eddington's anecdote

PROCEEDINGS
OF THE
NATIONAL ACADEMY OF SCIENCES

Volume 18

March 15, 1932

Number 3

ON THE RELATION BETWEEN THE EXPANSION AND THE
MEAN DENSITY OF THE UNIVERSE

BY A. EINSTEIN AND W. DE SITTER

Communicated by the Mount Wilson Observatory, January 25, 1932

In a recent note in the *Göttinger Nachrichten*, Dr. O. Heckmann has pointed out that the non-static solutions of the field equations of the general theory of relativity with constant density do not necessarily imply a positive curvature of three-dimensional space, but that this curvature may also be negative or zero.

Einstein-de Sitter model revisited

Über das sogenannte kosmologische Problem.

Wenn wir Raum und Zeit von relativistischer Physik absolut nennen, so hat dies folgende Bedeutung. Erstens hat der Raum und Zeit ein bestimmtes, ohne die Bedeutung von einer Realität von dem die Masse. Die Koordinaten hängen auf das gewählte Bezugssystem ab. Zweitens bedeutet die Relativität der Zeit, dass die Bedeutung von physikalischen Behauptungen haben, die nicht oder falsch sein können. Das dritte System bedeutet eine Realität und eine Welt in der Tätigkeitszeit einget. Zweitens ist das physikalisch Real, was mit dem Raum & Zeit bezeichnet wird, in einem Geometrisystem unabhängig von dem Verhalten des übrigen physikalisch-Realen, d. h. unabhängig von den Körpern. Im Hinblick die Beziehungen zwischen Massenanteilen, die alles in der Welt der Materie und Masse zu gewinnen sind, ist nachfolgend über die von der Leistung und Lösung der Körper unabhängig, ebenso das Interaktion. In Physik Raum ist gewissermaßen physikalisch vorhanden aber nicht physikalisch beeinflusst.

■ Einstein's cosmology review of 1933

Overview of dynamic models from first principles

Significant exposition

Culminates in Einstein-de Sitter model

Including time evolution

$$2\Lambda \frac{d^2 A}{dt^2} + \left(\frac{dA}{dt}\right)^2 = 0$$

$$3 \left(\frac{dA}{A}\right)^2 = \kappa \rho c^2$$

$$3h^2 = \kappa \rho c^2 (= 8\pi K \rho)$$

$$\Lambda = c(t - t_0)^{\frac{2}{3}}$$



■ Published in 1933

French book (transl. Solovine); small print run

Intended for scientific journal; not submitted

SUR LA STRUCTURE COSMOLOGIQUE DE L'ESPACE ⁽¹⁾

Si nous appelons l'espace et le temps de la physique prérelativiste « absolus », il faut y voir la signification suivante. Tout d'abord l'espace et le temps et, par suite, le système de référence, y figurent dans le même sens comme réalité que, par exemple, la masse. Les coordonnées du système de référence choisi y correspondent immédiatement à des résultats de mesure (?). Les propositions de géométrie et de cinématique signifient pour cette raison des relations entre des mesures ayant la valeur d'affirmations physiques, qui peuvent être vraies ou fausses. Le système d'inertie possède une réalité physique, parce que son choix entre dans la loi d'inertie. En second lieu, cette réalité physique, qui est désignée par les termes espace + temps, est, quant à ses lois, indépendante du comportement des autres réalités physiques, par exemple, des corps.

■ Parameters extracted

Critical density of 10⁻²⁸ g/cm³ : reasonable

Timespan of 10¹⁰ years: incorrect estimate

Conflict with astrophysics: attributed to simplifications

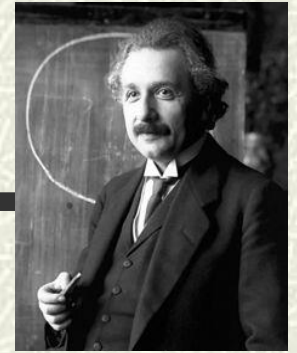
Some key quotes (Einstein 1933)

“Since, according to the general theory of relativity, the metric properties of space are not given in themselves but are instead determined by material objects that force a non-Euclidean character on the continuum, a problem arises....since we may assume that the stars are distributed with a finite density everywhere in the world, that is, a non-zero average density of matter in general, there arises the question of the influence of this mean density on the (metric) structure of space on a large scale; this is the so-called cosmological problem “

“Thus the theory can now, without the introduction of a λ -term, accommodate a finite (mean) density of matter ρ on the basis of equations (1), using relation 3(a) with P (and ρ) variable over time”

“It follows from these considerations that in the light of our present knowledge, the fact of a non-zero density of matter need not be reconciled with a curvature of space, but instead with an expansion of space. Of course, this does not mean that such a curvature (positive or negative) does not exist. However, there is at present no indication of its existence. In any case, it may well be substantially smaller than might have been suggested by the original theory (see equation 5).”

Einstein's cosmology: conclusions



Major test for general relativity

Conscious of assumptions of homogeneity, isotropy

Embraces dynamic cosmology

New evidence – new models (JMK)

Timespan of expanding models puzzling

Steady-state universe?



Hubble constant revised

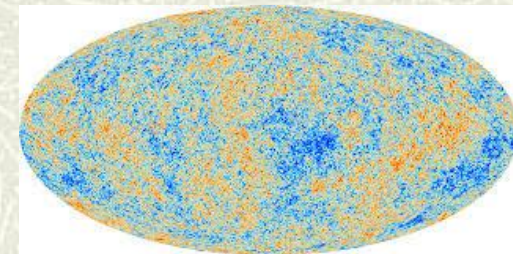
Evolving models (less contrived)

Cosmic constant not necessary

Extraction of parameters compatible with observation

Closed and open models

Timespan problem attributed to simplifying assumptions



Cosmic microwave background
Homogeneous, flat universe

Verdict (1933, 1945): more observational data needed

No mention of origins

Einstein's greatest hits (cosmology)



Einstein's model of the Static Universe (1917)

First relativistic model of the cosmos

Einstein's steady-state model (Jan 31)

Natural successor to static model: abandoned

Friedman-Einstein model of the Universe (1931)

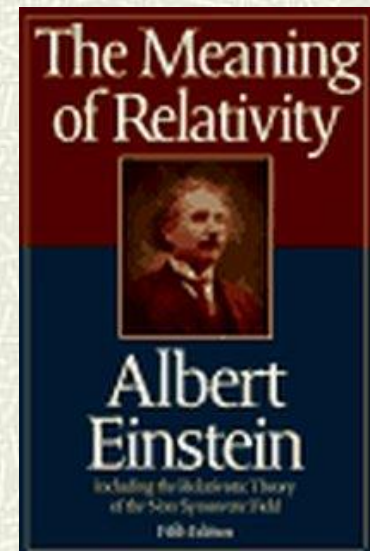
Use of Hubble constant to extract observational parameters

Einstein-de Sitter model of the Universe (1932)

1933 review: 1945 review (Appendix)

Conversations with Gamow, Godel, Straus

No mention of origins



Further reading

The European Physical Journal volume 39 · number 1 · February 2014

EPJ H

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Historical Perspectives on Contemporary Physics

An image of the blackboard used in Einstein's 2nd Rhodes lecture at Oxford in April 1931 (reproduced by permission of the Museum of the History of Science, University of Oxford)

$$D = \frac{1}{c} \frac{1}{l} \frac{dl}{dt} = \frac{1}{c} \frac{1}{P} \frac{dP}{dt}$$

$$D^2 = \frac{1}{P^2} \frac{P_0 - P}{P} \sim \frac{1}{P^2} \quad (1a)$$

$$D^2 = \frac{k_0}{3} \frac{P_0 - P}{P} \sim \frac{1}{3} k_0 \quad (2a)$$

$$D^2 \sim 10^{-53}$$

$$\rho \sim 10^{-26}$$

$$P \sim 10^8 \text{ g/cm}^3$$

$$t \sim 10^{10} (10^{11}) \text{ s}$$

Einstein's cosmic model of 1931 revisited: An analysis and translation of a forgotten model of the universe

by Cormac O'Riadafeartaigh and Brendan McCann

edp sciences Springer

Eur. Phys. J. H
DOI: 10.1140/epjh/e2014-50011-x

THE EUROPEAN PHYSICAL JOURNAL H

Einstein's steady-state theory: an abandoned model of the cosmos

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Abstract. We present a translation and analysis of an unpublished manuscript by Albert Einstein in which he attempted to construct a 'steady-state' model of the universe. The manuscript, which appears to have been written in early 1931, demonstrates that Einstein once considered a cosmic model in which the mean density of matter in an expanding universe is maintained constant by the continuous formation of matter from empty space. This model is very different to previously



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Historical Perspectives on Contemporary Physics

Cover of the French translation of three Einstein articles by Maurice Solovine, published in 1933.

From: *Einstein's cosmology review of 1933: A new perspective on the Einstein-de Sitter model of the cosmos*

by Cormac O'Riadafeartaigh, Michael O'Keefe, Werner Nahm and Simon Mitton

edp sciences Springer

Coda: Einstein vs Hoyle

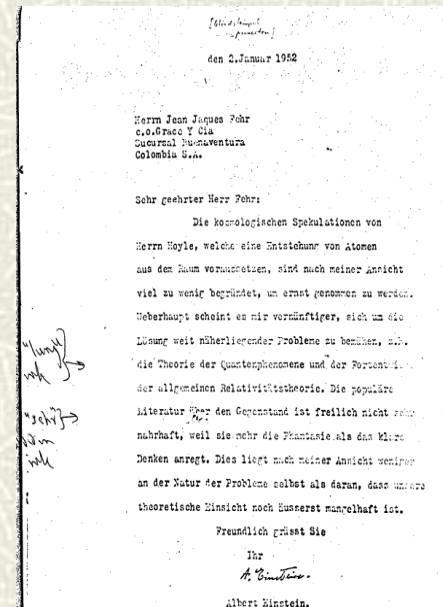
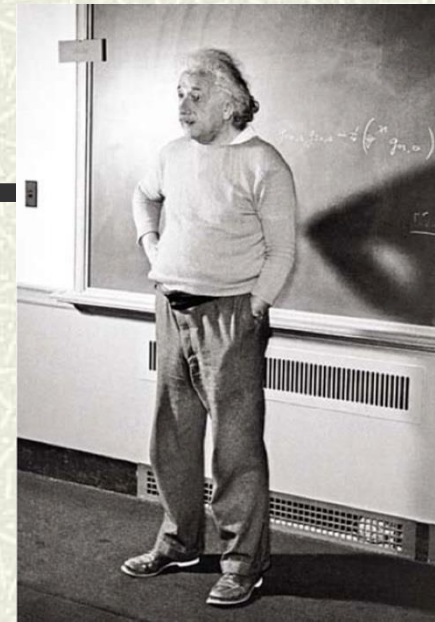
Fred Hoyle in Princeton (1952, 53)

Einstein remark to Manfred Clynes

“Romantic speculation” (Michelmore 1962)

Letter to Jean Jacques Fehr (1952)

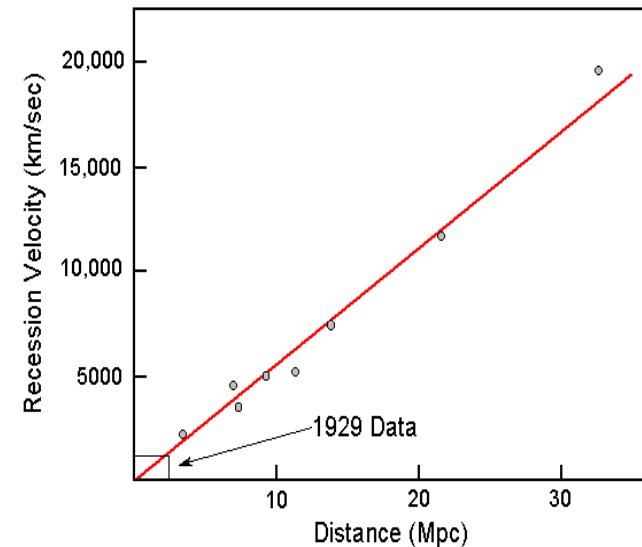
“The cosmological speculations of Mr Hoyle, which presume a formation of atoms from space, are in my view much too poorly grounded to be taken seriously. On the whole, it seems to me more reasonable to seek a solution to problems far closer to hand, e.g., the theory of quantum phenomena or the further development of the general theory of relativity. The popular literature on the subject is not very fruitful, as it encourages flights of fancy rather than clear thinking. In my opinion, this is less because of the nature of the problem itself than because our theoretical insight is still extremely deficient.”



Observational parameters needed (1930s)

- # **Spatial curvature** $k = -1, 0, 1?$
- # **Cosmic constant** $\lambda = 0?$
- # **Deceleration** $q_0 = -\ddot{R}/\dot{R}^2$
- # **Density of matter** $\rho < \rho_{crit} ?$
- # **Timespan** $\tau = 10^{10} \text{ yr}?$
- # **Hubble constant** $\dot{R}/R = 500 \text{ kms}^{-1}\text{Mpc}^{-1}?$

Hubble & Humason (1931)



*What do redshifts represent?
Is expansion a local effect?*

Hubble and Tolman 1935

Einstein's steady-state model and cosmology today

Dark energy (1998)

Accelerated expansion (observation)

Positive cosmological constant

Einstein's dark energy

“The conservation law is preserved in that, by setting the λ -term, space itself is not empty of energy; its validity is well known to be guaranteed by equations (1).”

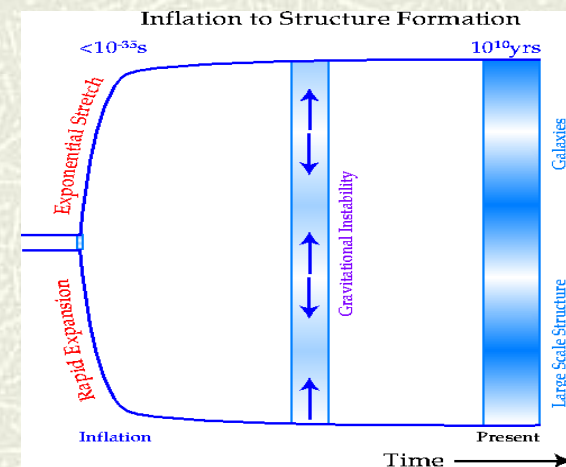
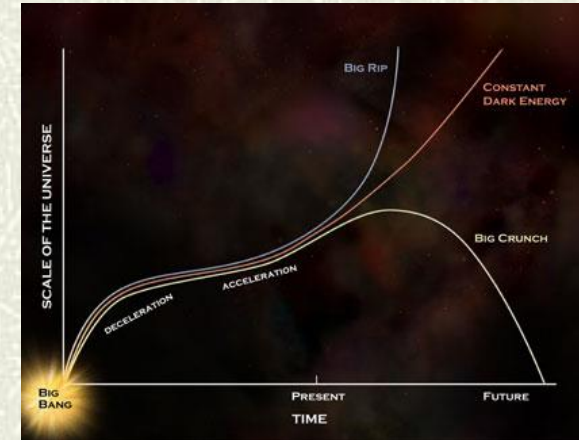
Cosmic inflation

Inflationary models use de Sitter metric

Used in all steady-state models

Flat curvature, constant rate of matter creation

Different time-frame!



Einstein vs Hoyle

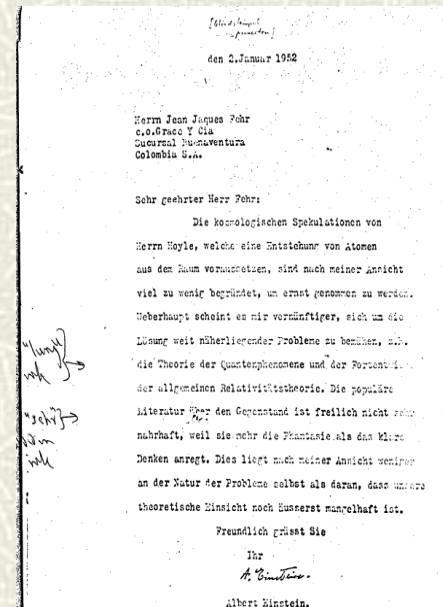
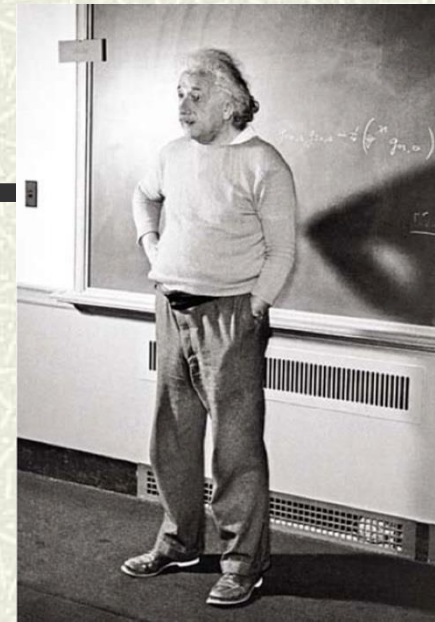
Hoyle in Princeton (1952, 53)

Einstein remark to Manfred Clynes

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Letter to Jean Jacques Fehr (1952)

- # *“The cosmological speculations of Mr Hoyle, which presume a formation of atoms from space, are in my view much too poorly grounded to be taken seriously. On the whole, it seems to me more reasonable to seek a solution to problems far closer to hand, e.g., the theory of quantum phenomena or the further development of the general theory of relativity. The popular literature on the subject is not very fruitful, as it encourages flights of fancy rather than clear thinking. In my opinion, this is less because of the nature of the problem itself than because our theoretical insight is still extremely deficient.”*



Steady-state cosmology today

Observable universe not in a steady state

Evolution of galaxies

Cosmic microwave background

Inflationary cosmology = steady-state model

de Sitter metric

Steady-state model with different time-frame! (Hoyle 1990)

Matter creation term not mandatory in Hoyle models (McCrea 1951)

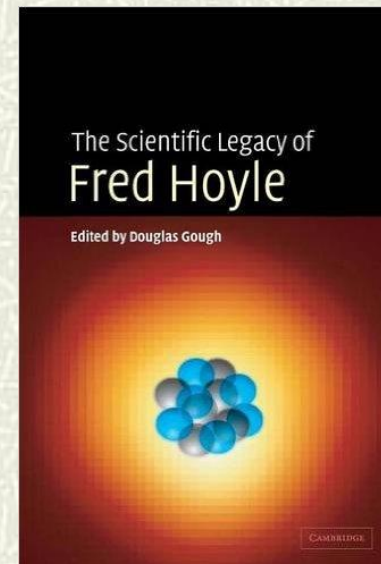
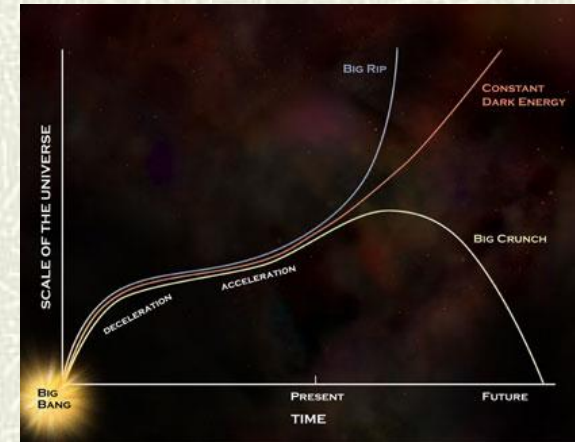
Eternal inflation

Different regions undergo different inflation?

Inflation begets further inflation (Vilenkin 1983; Linde 1986)

Observable universe embedded in global steady-state cosmos?

Hoyle's revenge! (Hoyle and Narlikar 1966; Barrow 2005)



Einstein's steady-state model and cosmology today

Accelerated expansion (1998)

Supernova measurements

Dark energy – positive cosmological constant



Einstein's dark energy

“The conservation law is preserved in that, by setting the λ -term, space itself is not empty of energy; its validity is well known to be guaranteed by equations (1).”

Anticipates positive cosmological constant

De Sitter line element

$$ds^2 = - e^{at} (dx_1^2 + dx_2^2 + dx_3^2) + c^2 dt^2 \dots$$

Necessary for all steady-state models

Identical to inflationary models (different time-frame)

Einstein's cosmology (traditional view)

Static, bounded model of the cosmos (1917)

Resistance to empty de Sitter 'static' solution

Resistance to time-varying models

Friedman, Lemaitre

Einstein the conservative

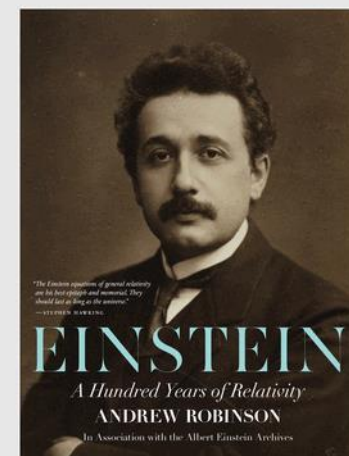
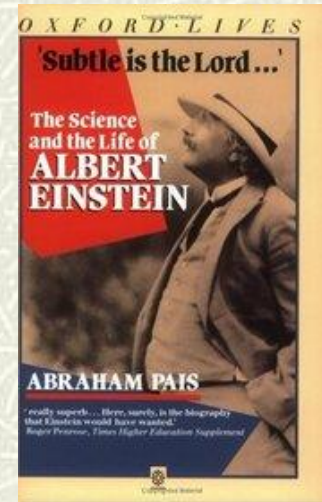
Hidebound by philosophical prejudice

Disinterested in evolving cosmologies

Dynamic model 1932 (1931)

Minimalist, cursory models

A new perspective? Guided by physical intuition



Taking $T_{44} = \rho c^2$ (all other components zero) in the *time* component of equation (1) we obtain $\left(R_{44} - \frac{1}{2}g_{44}R\right) - \lambda g_{44} = \kappa\rho c^2$.

This gives on analysis $-3\alpha^2/4 + 3\alpha^2/2 - \lambda c^2 = \kappa\rho c^2$
the second of Einstein's simultaneous equations.

From the *spatial* component of equation (1), we obtain $\left(R_{ii} - \frac{1}{2}g_{ii}R\right) - \lambda g_{ii} = 0$.

This gives on analysis $3\alpha^2/4 - 3\alpha^2/2 + \lambda c^2 = 0$
for the first of the simultaneous equations.

It is plausible that Einstein made a sign error here, initially getting $3\alpha^2/4 + 3\alpha^2/2 + \lambda c^2 = 0$ for this equation.